

Summary Files

Summary files have the filename 'summary.txt'. They are text (ASCII) files containing one line for each time step. Each line is divided into 23 columns, containing the following data:

| Column Number | Datum | Description |
|---------------|--------------------------|--|
| 1 | i | Step number |
| 2 | t | Age (years) |
| 3 | M | Mass (M_{\odot}) |
| 4 | $\text{Log}_{10} L$ | Luminosity (L_{\odot}) |
| 5 | $\text{Log}_{10} R$ | Radius (R_{\odot}) |
| 6 | $\text{Log}_{10} T_s$ | Surface temperature (K) |
| 7 | $\text{Log}_{10} T_c$ | Central temperature (K) |
| 8 | $\text{Log}_{10} \rho_c$ | Central density (kg m^{-3}) |
| 9 | $\text{Log}_{10} P_c$ | Central pressure (N m^{-2}) |
| 10 | ψ_c | Central electron degeneracy parameter |
| 11 | X_c | Central hydrogen mass fraction |
| 12 | Y_c | Central helium mass fraction |
| 13 | $X_{C,c}$ | Central carbon mass fraction |
| 14 | $X_{N,c}$ | Central nitrogen mass fraction |
| 15 | $X_{O,c}$ | Central oxygen mass fraction |
| 16 | T_{dyn} | Dynamical timescale (years) |
| 17 | T_{KH} | Kelvin-Helmholtz timescale (years) |
| 18 | T_{nuc} | Nuclear timescale (years) |
| 19 | L_{pp} | Luminosity from PP chain (L_{\odot}) |
| 20 | L_{CNO} | Luminosity from CNO cycle (L_{\odot}) |
| 21 | $L_{3\alpha}$ | Luminosity from triple-alpha reactions (L_{\odot}) |
| 22 | L_Z | Luminosity from metal burning (L_{\odot}) |
| 23 | L_{ν} | Luminosity of neutrino losses (L_{\odot}) |
| 24 | M_{He} | Mass of helium core (M_{\odot}) |
| 25 | M_C | Mass of carbon core (M_{\odot}) |
| 26 | M_O | Mass of oxygen core (M_{\odot}) |
| 27 | R_{He} | Radius of helium core (R_{\odot}) |
| 28 | R_C | Radius of carbon core (R_{\odot}) |
| 29 | R_O | Radius of oxygen core (R_{\odot}) |

How close to degeneracy are we? (White Dwarfs and Neutron Stars are made of degenerate electrons and neutrons, respectively)

$t_{\text{dyn}} \sim (\text{sound speed})/(\text{length scale})$

How long it would take star to radiate away all its gravitational energy if nothing opposed its collapse

How long until the star burns all of whatever it is currently burning?

Luminosity from He burning (if star is massive enough to eventually burn He)

Each nuclear process loses neutrinos - how much energy do they carry away per second?